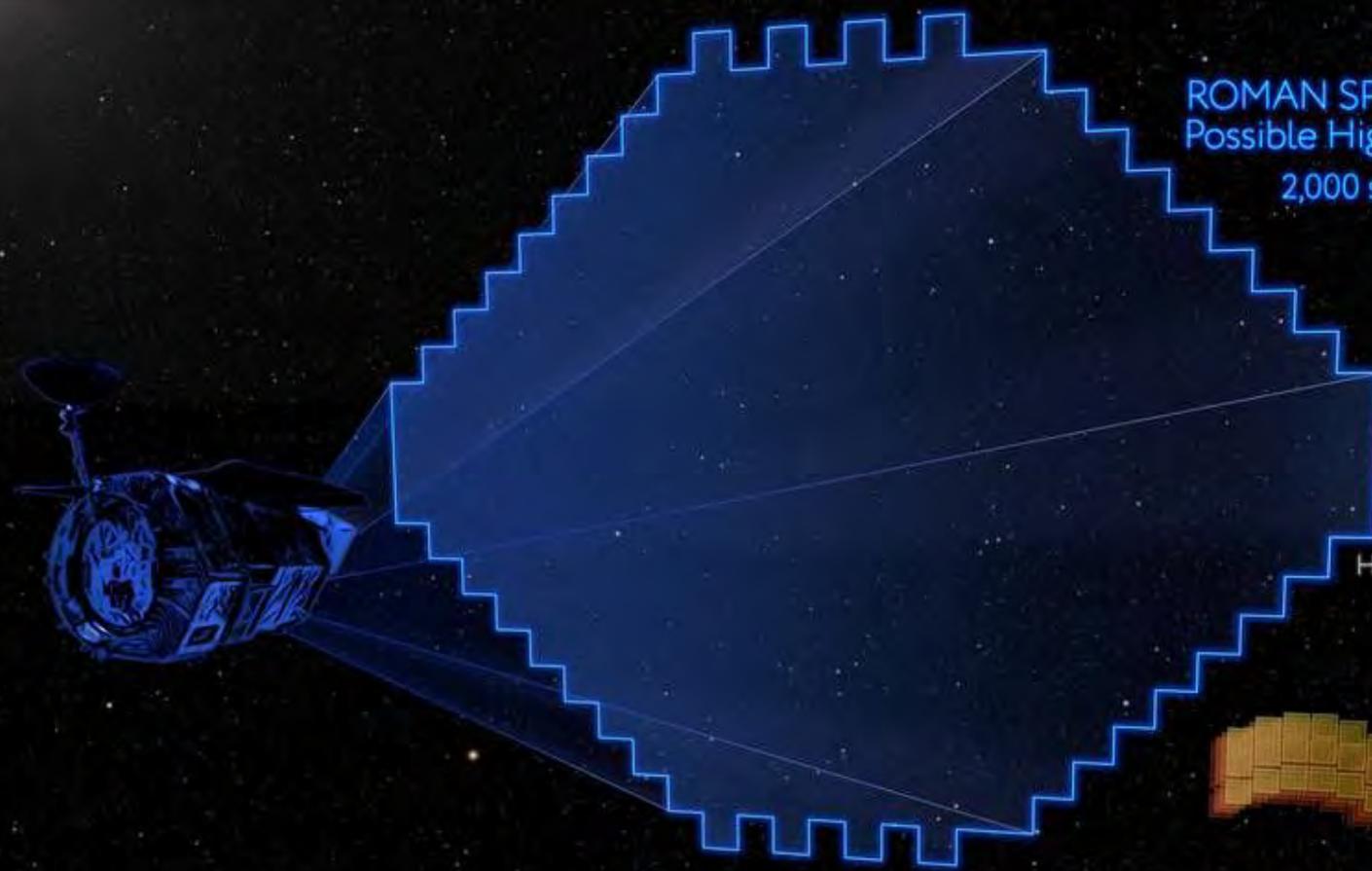


# MESA Isochrones and Stellar Tracks (MIST)

Nancy Grace Roman Space Telescope

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A diagram showing the Roman Space Telescope on the left, emitting a large, blue, stepped cone of light that represents its survey area. The cone is wider than the Hubble Space Telescope's survey area, which is shown as a smaller, similar cone below it. The background is a dark field of stars.

ROMAN SPACE TELESCOPE  
Possible High Latitude Survey  
2,000 square degrees

HUBBLE SPACE TELESCOPE  
COSMOS Program  
1.6 square degrees

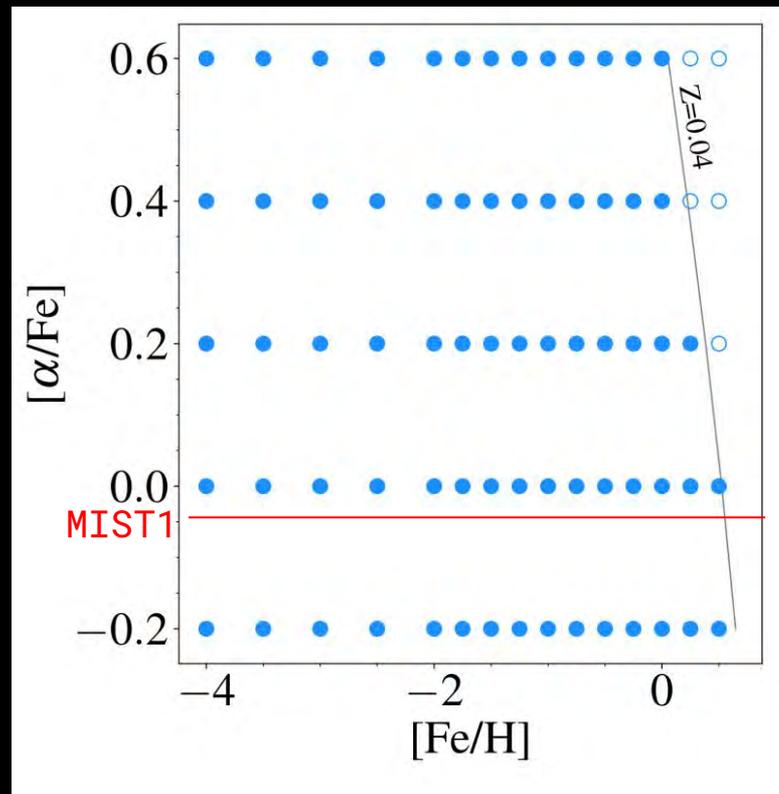


Each Roman tile will be four  
stacked and offset images in  
different infrared wavelengths

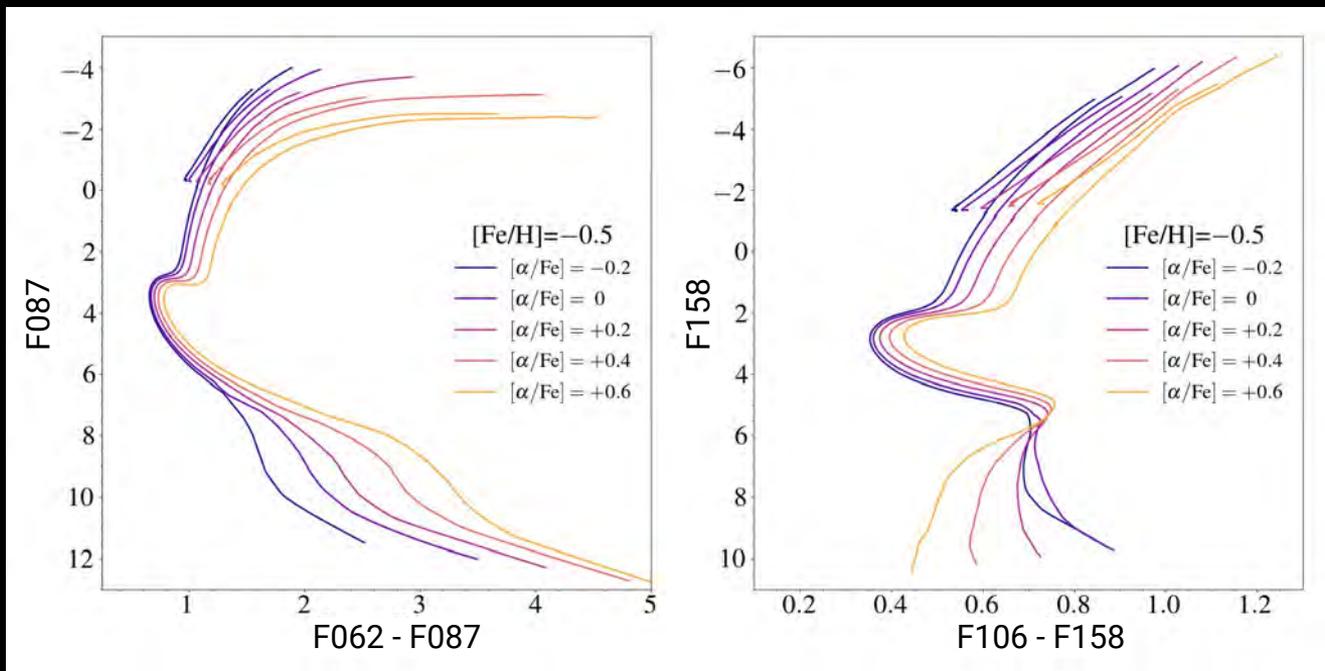
# What is MIST? Why is it important? What is new?

The MESA Isochrones and Stellar Tracks (MIST) project uses the Modules for Experiments in Stellar Astrophysics (MESA) Software Instrument to “survey” a wide range of stars, from M dwarfs to very massive stars ( $0.1$  to  $300 M_{\odot}$ ).

SIT funding supports the development of MIST2, enhancing the science yield of Roman per Astro2020 Decadal Survey.

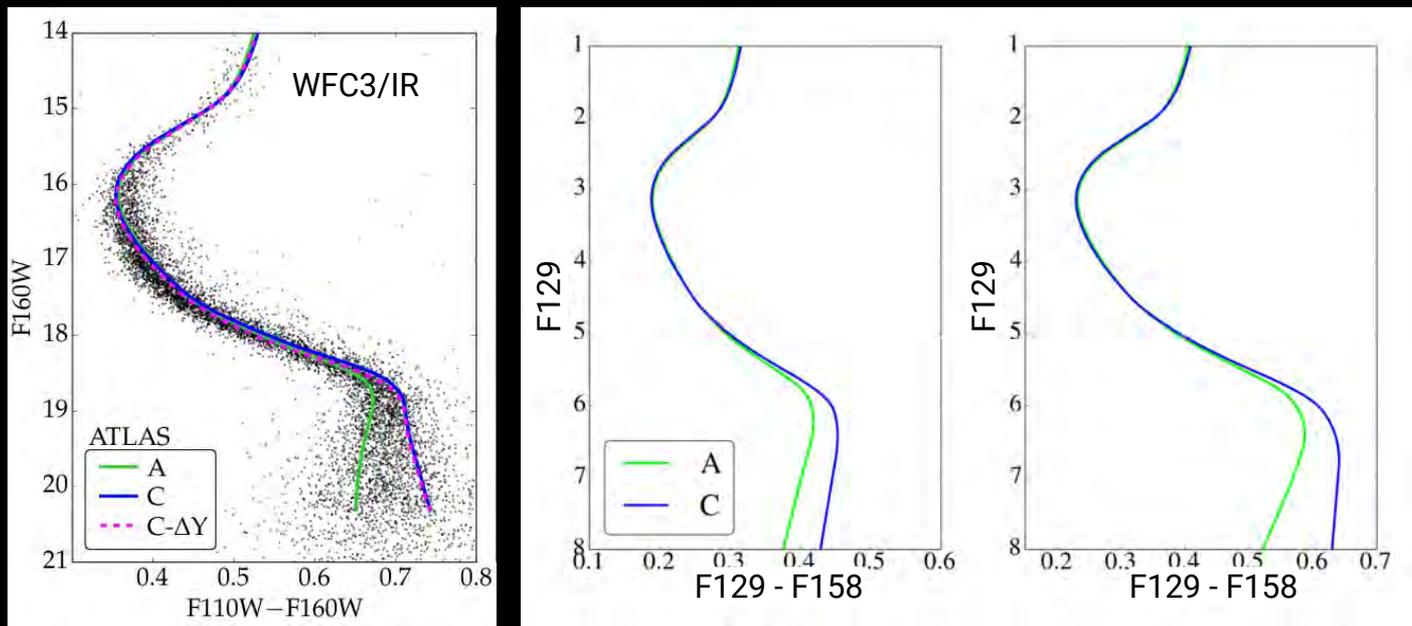


# MIST Development for Roman



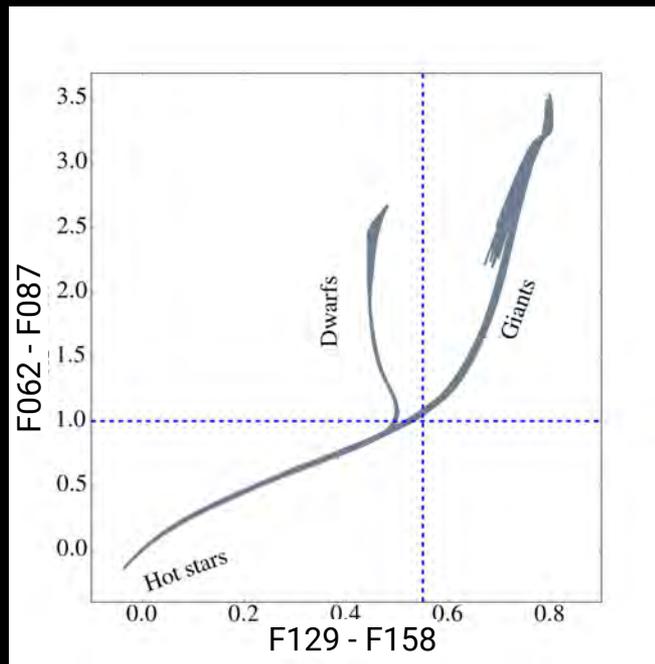
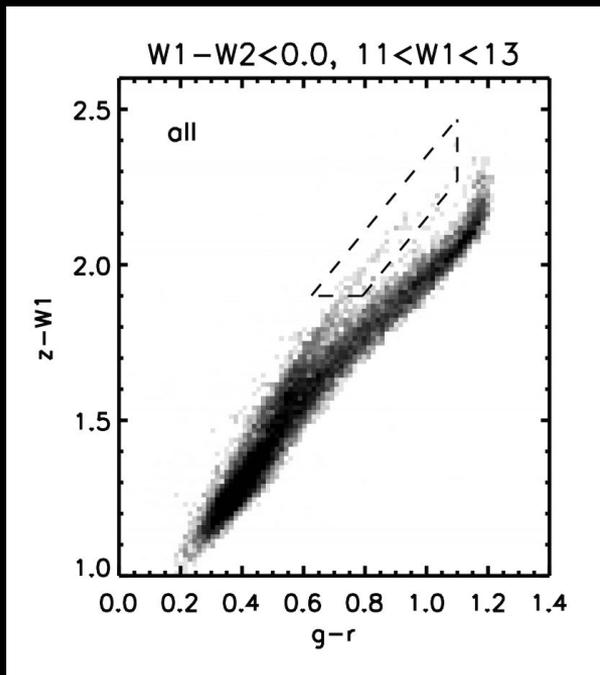
MIST has the potential to inform Roman science before the telescope reaches orbit, by showing the power of different filter combinations. (Based on preliminary WFI filter curves.) Assess possibilities, translate results between other telescopes and Roman.

# MIST & Roman: Multiple populations in GCs



Most, if not all, globular clusters harbor multiple stellar populations. These are characterized by, among other things, systematic variations in the C,N,O abundances. Near-IR H<sub>2</sub>O lines enable us to study MPs with photometry, as shown here in HST data of NGC 6752 (Dotter et al. 2015). Roman WFI filters J129, H158, and F184 prove very useful!

# MIST & Roman: Dwarf-Giant Separation



Conroy et al. (2018) show the power of large photometric surveys to derive dwarf-giant separation cuts. Such techniques have been used extensively to find elusive stellar populations. Here we show that Roman WFI photometry can be utilized in a similar way, homogeneously.

# MIST & Roman: Next Steps

What's next for MIST+Roman, beyond v2?

- Better treatment of AGB stars, dominant source of light in the near-IR
- Spot model for cool dwarfs on the lower main sequence
- Expanded support for different extinction treatments:

MIST1 uses Cardelli et al. (1989) with  $R_v=3.1$

MIST2 uses Fitzpatrick & Massa (2007) with  $2 \leq R_v \leq 5$

... but other, more sophisticated treatments exist

